

PN-CARBON YIELDS AND THE CHEMICAL EVOLUTION OF THE GALAXY

Leticia Carigi

Instituto de Astronomía, Universidad Nacional Autónoma de México

RESUMEN

Calculo dos conjuntos de rendimientos estelares observacionales de carbono para estrellas de masa intermedia y baja basados en abundancias derivadas a partir de líneas permitidas y prohibidas de C II y C III ($\lambda 4267$ y $\lambda \lambda 1906 + 1909$, respectivamente). Comparando los valores de C/O observados en la vecindad solar y en regiones H II Galácticas con aquellos predichos por modelos de evolución química para la Galaxia que suponen estos rendimientos observacionales, se concluye que las abundancias de C basadas en líneas permitidas son mejores que aquellas derivadas a partir de líneas prohibidas.

ABSTRACT

Two sets of observational carbon stellar yields for low-and-intermediate mass stars are computed from planetary nebulae abundances derived from C II $\lambda 4267$ and C III $\lambda \lambda 1906 + 1909$ lines, respectively. By comparing C/O values observed in stars of the solar vicinity and Galactic H II regions with those predicted by chemical evolution models for the Galaxy, which assume these observational yields, I conclude that the C abundances derived from permitted lines are better than those derived from forbidden lines.

Key Words: **GALAXY: ABUNDANCES — GALAXY: EVOLUTION — PLANETARY NEBULAE: GENERAL**

1. INTRODUCTION

The $N(C^{++})/N(H^+)$ values derived from the permitted line (PL) C II $\lambda 4267$ are higher, by as much as a factor of 10, than those determined from the forbidden lines (FL) C III $\lambda \lambda 1906 + 1909$. Several explanations for this discrepancy have been presented in the literature, but the problem remains open. Since PNe are important for the C enrichment of the interstellar medium, a successful chemical evolution model for the solar vicinity and the Galactic disk (Carigi 2000) is used to discriminate between the PN-C abundances derived from permitted lines and the PN-C abundances obtained from forbidden lines.

2. PN OBSERVATIONAL YIELDS

Average $(C/H)_{PN}^{PL}$ and $(C/H)_{PN}^{FL}$ values are computed from 15 type I PNe (PNI), and from 21 type II and III PNe (PNII/III). $\langle (C/H)_{PNI}^{PL} \rangle$ is calculated from the $N(C^{++})/N(O^{++})$ and $N(O)/N(H)$ values given by Peimbert et al. (1995a). The $\langle (C/H)_{PNI}^{FL} \rangle$ value is obtained from $\langle (C/H)_{PNI}^{PL} \rangle$ and the average of $N(C^{++})/N(H^+)_{PL}/N(C^{++})/N(H^+)_{FL}$ ratios taken from Peimbert et al. (1995b). The $\langle (C/H)_{PNII/III}^{PL} \rangle$ and $\langle (C/H)_{PNII/III}^{FL} \rangle$ values are computed from the $N(C^{++})/N(H^+)$ given by Peimbert et al. (1995b) and corrected for the contribution of $N(C^+)/N(H^+)$ by adding 0.1 dex.

According to stellar evolution models, the PNI progenitors are stars with initial mass between 2.4 and $8 M_{\odot}$ and the PNII/III progenitors are stars with $0.8 < m/M_{\odot} < 2.4$. The C yields for PN progenitors (C_{PN} yields) are calculated based on the $\langle (C/H)_{PN} \rangle$ from permitted lines and forbidden lines, neglecting the ejected mass by winds and assuming that the $\langle (C/H)_{PN} \rangle$ values are independent of the initial metallicity of the

progenitors. I assume that the O yields for PN progenitors are null. The mass ejected by PN progenitors are taken from van den Hoek & Groenewegen (1997).

3. MODELS

All models are built to reproduce the observed gas fraction distribution of the Galaxy and the observed O/H Galactic gradient from 4 to 10 kpc. Models adjust the rise of C/O with metallicity in the solar neighborhood shown by dwarf stars, the Sun, and Orion; and the decrease of the C/O abundance with Galactocentric distance derived from the Galactic H II regions M17, M8, and Orion.

The models are very similar to the infall model of Carigi (2000), but in this work there are some differences in the assumptions about stellar yields:

a) Just one set of metal-dependent stellar yields from massive stars ($8 < m/M_{\odot} < 120$) is considered: Geneva yields (Maeder 1992). Carigi concluded that only models with the Geneva yields can reproduce both, the increase of C/O with Z in the solar vicinity and the negative C/O gradient.

b) Four sets of stellar yields for low and intermediate mass stars ($0.8 < m/M_{\odot} < 8$) are used: i) two metal-dependent-theoretical ones: van den Hoek & Groenewegen (1997, Amsterdam yields), and Marigo et al. (1996, 1998) and Portinari et al. (1998) (Padova yields); ii) two metal-independent-observational stellar yields: PN yields from permitted lines and PN yields from forbidden lines.

C/O depends on the initial mass function and the C and O yields. Since the initial mass function and the massive-star yields are fixed, the C/O value is used to test the different sets of C and O yields for low and intermediate mass stars

4. RESULTS AND CONCLUSIONS

By comparing observed and theoretical C yields, it can be noted that: i) there is a very good agreement between the $C_{\text{PN}}^{\text{PL}}$ yields and those predicted by stellar evolution models with $Z = Z_{\odot}$, and ii) the $C_{\text{PN}}^{\text{FL}}$ yields are lower, by as much as a factor of 4, than those predicted by stellar evolution models with $Z = Z_{\odot}$.

From chemical evolution models of the Galaxy, I conclude that:

a) Models with C_{PN} yields derived from permitted lines match all the observational constraints, in particular they reproduce the C/O absolute values observed in stars of the solar vicinity and in H II regions of the Galactic disk.

b) Models with C_{PN} yields from forbidden lines fail to reproduce the C/O ratios in dwarf stars of different ages in the solar vicinity, the Sun, and M17.

c) Models with C_{PN} yields derived from permitted lines agree with models based on theoretical yields, in particular showing better agreement with models based on the Padova yields than models based on the Amsterdam yields.

d) The C/O values predicted with C_{PN} yields derived from permitted lines are 0.1 dex higher than those obtained with the Amsterdam yields, and 0.05 dex lower than those computed with the Padova yields.

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Leticia Carigi: Instituto de Astronomía, UNAM, Apartado Postal 70-264, 04510 México D.F., México
 (carigi@astroscu.unam.mx)